

Photometry and Light Curve Modeling of HO Piscium and V535 Pegasi

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Abstract In this article we will present the photometric study of the overcontact binaries HO Psc and V535 Peg. The data were acquired with the 304-mm telescope of RIAAM Observatory, and after the data reduction and photometry, the main parameters of the systems such as temperatures, inclination, and mass ratio were found using modeling in PHOEBE software.

1. Introduction

Studies of eclipsing binary stars are currently of interest because of testing models and understanding their various intrinsic properties (Terrell 2006). HO Psc (Martignoni 2006) and V535 Peg (Geske *et al.* 2006) are also overcontact eclipsing binaries, so they share a common envelope of material. W UMa system light curves usually have equal depth of primary and secondary minima, that is because both components have almost equal temperature. In W UMa variables, usually the components are so close that gravitational effects causes deformations of components. Information on these stars can be seen in Figure 1 and Table 1.

2. Observation and data reduction

We used the Research Institute for Astronomy and Astrophysics of Maragheh (RIAAM) observatory equipment which include a 304.8mm schmidt-cassegrain telescope and a SBIG STX-16803 CCD. The data were captured from July 2016 to October 2017 in BVR filters. The telescope was guided with a DMK31AU03 CCD mounted on a small telescope with focal length of 1000mm. We also used 2×2 binning, and the CCD's temperature was fixed on -35°C with 75% of cooling fan power. The IRAF package was used for reducing the bias and dark frames and also dividing the flat field frame, which we

captured in twilight from sky horizon at opposite direction of the sunrise. The reduced data were used for aperture photometry, so the photometry files were made in columns of HJD, object magnitude, and check star magnitude for each filter.

All of the magnitude points of the check star were subtracted from the average of them in order to find the variation range of the magnitude in observing time, and then they were subtracted from the object's magnitude (as seen in Tables 6 and 7). The HJD were converted to orbital phase using Equation 1 and the light curves have been plotted as seen in Figures 2 and 3.

$$\text{Phase} = \text{decimal}[(\text{HJD}_0 - \text{Epoch}) / \text{period}] \quad (1)$$

The data of period of 0.4 to 0.6 and 0.9 to 1.1 for each minima were exported to an ASCII file in HJD and magnitude columns in order to calculate the minima time. The table curve software was used for peak fitting the data and calculating the minima time and error as seen in Table 2, and the O-C diagrams were plotted (Figure 4).

The data for HO Psc and V535 Peg are given in Tables 6 and 7, respectively, at the end of this article.

3. Light curve modelling

In order to find the physical parameters of the systems we tried to achieve the best fitted model using the PHOEBE software (Prša and Zwitter 2005), which uses the Wilson-Devinney code (WD; Wilson and Devinney 1971). The data were imported as an ASCII file with columns of phase and magnitude of the objects. The initial values of the temperatures of the systems were used considering their color indexes. As we have close binaries, the difference between the surface temperatures of the components was almost equal.

The next important parameter in PHOEBE is the mass ratio (q) of the systems. We tried about 20 steps for initial value of the mass ratio from 0.4 to 3, and considering the shape of the synthetic light curve and the χ^2 value, $q = 0.9$ for HO Psc and $q = 0.6$ for V535 Peg were used.

The limb darkening values have been added from the van Hamme (1993) table with logarithmic law for the stars with

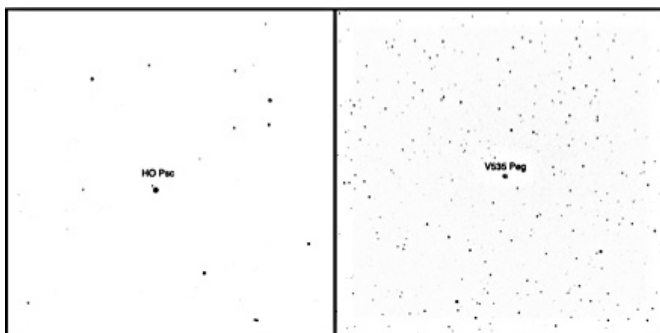


Figure 1. Field of view of the objects, left: HO Psc, right: V535 Peg.

Table 1. Objects.

Object	R.A. (2000)			Dec. (2000)		B	Magnitude (simbad)				
	h	m	s	deg	'		V	J	H	K	
HO Psc	01	30	16.466	+13	33	25.08	11.0	11.50	9.659	9.29	9.215
V535 Peg	22	36	16.7640	+33	18	56.761	11.18	10.851	9.157	8.793	8.694

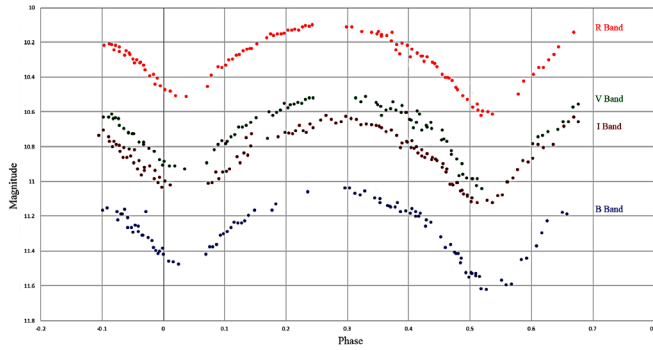


Figure 2. V535 Peg BVRI light curve.

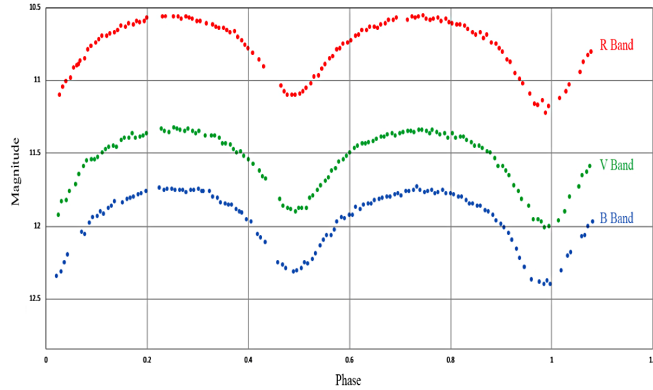


Figure 3. HO Psc BVR light curve.

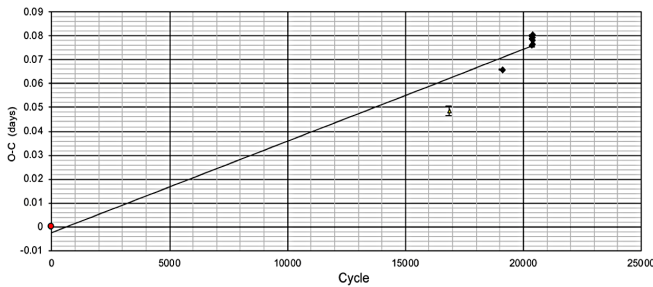
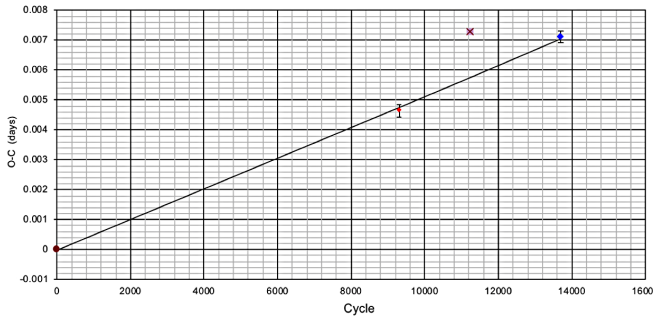


Figure 4. O-C diagrams for HO Psc (upper plot) and V535 Peg (lower plot).

$T_{\text{eff}} < 9000$ with convective layer (Al-Naimiy 1978). The gravity darkening values were also used from the table of Lucy (1967). After that, we tried to achieve the best physical parameters of inclination and temperature. With a lot of iterations, we used the calculation of PHOEBE to get the best fitted model considering the χ^2 without spot as seen in Table 3. In both systems' light curves, a difference in the magnitude was seen in the out-of-eclipse region that is known as the O'Connell effect (O'Connell

Table 2. Observed minima.

Object	Primary Minima	rms Error	Secondary Minima	rms Error
HO Psc	2457702.4888	$4.1e^{-5}$	2457702.3262	$0.9e^{-5}$
V535 Peg	2458035.3140	$6.32e^{-5}$	2458035.4725	$1.83e^{-5}$
	2458038.2224	$1.22e^{-5}$	2458038.383	$3.91e^{-5}$
	2458039.189	$2.94e^{-5}$	2458039.349	$1.39e^{-5}$

Table 3. Physical parameters.

Parameter	HO Psc	V535 Peg	Error
Period (days)	0.324747736	0.323003849	—
New epoch	2457702.4872	2458039.18664	—
Ω_1	3.35	2.85	0.03
Ω_2	3.35	2.85	0.03
q_{ptm}	0.90	0.58	0.01
Inclination	75.14°	72.56°	0.1
Limb Darkening (linear)	$x1 = 0.68$ $y1 = 0.18$	$x1 = 0.67$ $y1 = 0.21$	—
Limb Darkening (non-linear)	$x2 = 0.68$ $y2 = 0.18$	$x2 = 0.67$ $y2 = 0.20$	—
Gravity Darkening	$g1 = 0.5$	$g1 = 0.32$	—
	$g2 = 0.82$	$g2 = 0.32$	—
T_{eff1}	6674K	6730K	12
T_{eff2}	6228K	6509K	20
$L_1 (L_\odot)$	4.38	4.88	—
$L_2 (L_\odot)$	3.24	2.514	—
$R_1 (R_\odot)$	1.57	1.63	—
$R_2 (R_\odot)$	1.50	1.25	—
M_{bol1}	3.14	3.02	—
M_{bol2}	3.53	3.74	—
SMA	5.86	2.60	—

Table 4. HO Psc spot parameters.

	Colatitude	Longitude	Radius	Temperature
Primary Star	90	90	10	0.9

Table 5. V535 Peg spots parameters.

	Colatitude	Longitude	Radius	Temperature
Primary Star	90	90	20	0.7

1951), so we tried to add a spot (Tables 4 and 5). The fitted models of the systems are shown in Figures 5 and 6.

To test this model, we tried to calculate the luminosities and radii values using the empirical relationship between M_{bol} and T_{eff} given by Reed (1998) for the $T_{\text{eff}} < 9141$ as seen in Equation 2 and then, driving luminosity and radius with Equations 3 to 5.

$$\begin{aligned}
 BC = & -8.499 [\log(T) - 4]4 + 13.421 [\log(T) - 4]3 \\
 & - 8.131 [\log(T) - 4]2 - 3.901 [\log(T) - 4] - 0.438 \quad (2)
 \end{aligned}$$

$$M_{\text{bol}} = M_v + BC(T_{\text{eff}}) \quad (3)$$

$$M_{\text{bol}(*)} = M_{\text{bol}(\text{sun})} - 2.5 \log(L^* / L_{\text{sun}}) \quad (4)$$

$$R^2 = L / T_{\text{eff}}^4 \quad (5)$$

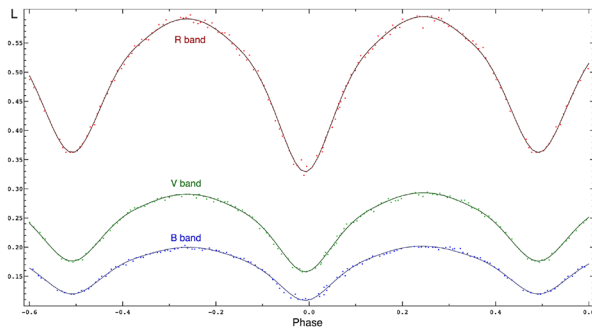


Figure 5. HO Psc fitted model.

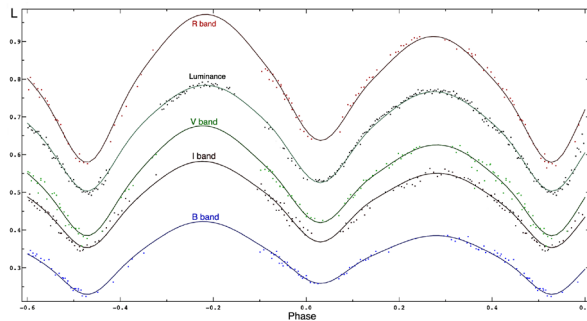


Figure 6. V535 Peg fitted model.

The measured values for radii and luminosities were the same as the values obtained in PHOEBE as seen in Table 3.

We also used the main sequence parameter table of Boyajian *et al.* (2013) for the mass of the components which were well matched with our stars, and considering the mass values we determined the semi-major axis values of the systems.

The 3D shapes of the systems were also drawn using BINARYMAKER software (Bradstreet and Steelman 2002; Figures 7 and 8).

4. Acknowledgement

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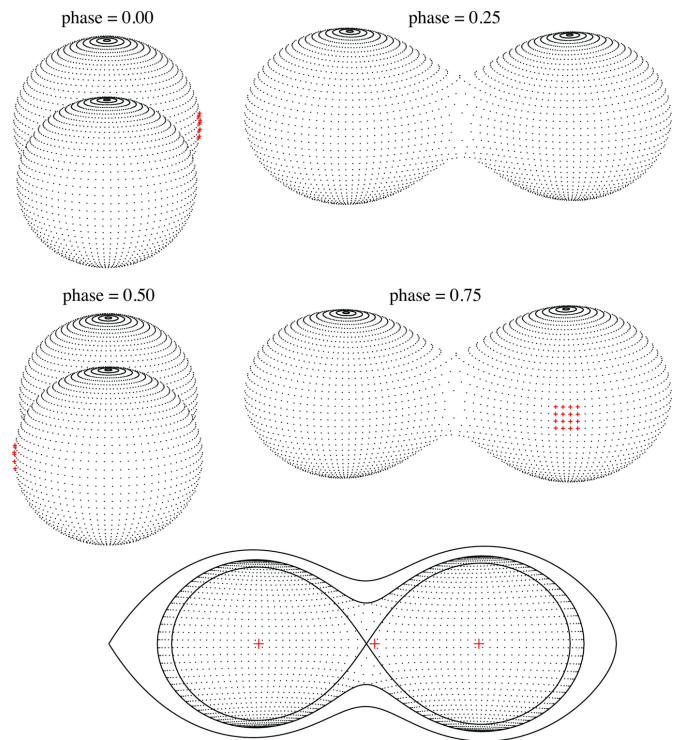


Figure 7. 3D shape of HO Psc.

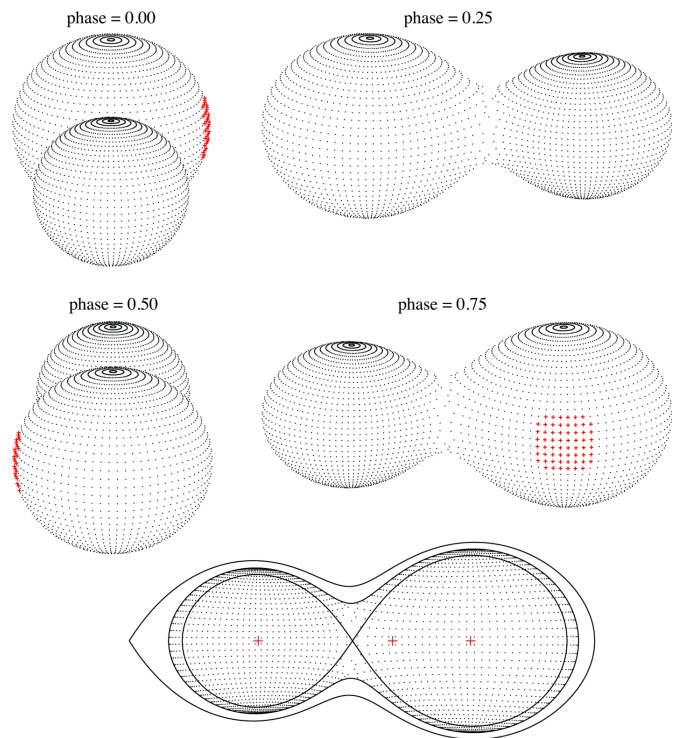


Figure 8. 3D shape of V535 Peg.

